

LETTER TO THE EDITOR

Dense gas without star formation: The kpc-sized turbulent molecular disk in 3C326 N ^{*}

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ABSTRACT

We report the discovery of a 3 kpc disk of few $10^9 M_\odot$ of dense, warm H_2 in the nearby radio galaxy 3C326 N, which shows no signs of on-going or recent star formation and falls a factor 60 below the Schmidt-Kennicutt law. VLT/SINFONI imaging spectroscopy shows broad (FWHM $\sim 500 \text{ km s}^{-1}$) ro-vibrational H_2 lines across all of the disk, with irregular profiles and line ratios consistent with shocks. The ratio of turbulent and gravitational energy suggests that the gas is highly turbulent and not gravitationally bound. In absence of the driving by the jet, short turbulent dissipation times suggest the gas should collapse rapidly and form stars, at odds with the recent star-formation history. Motivated by hydrodynamic models of rapid H_2 formation boosted by turbulent compression, we propose that the molecules formed from diffuse atomic gas in the turbulent jet cocoon. Since the gas is not self-gravitating, it cannot form molecular clouds or stars while the jet is active, and is likely to disperse and become atomic again after the nuclear activity ceases. We speculate that very low star-formation rates are to be expected under such conditions, provided that the large-scale turbulence sets the gas dynamics in molecular clouds. Our results illustrate that jets may create large molecular reservoirs as expected in 'positive feedback' scenarios of AGN-triggered star formation, but that this alone is not sufficient to trigger star formation.

Key words. Galaxies – ... – ...

1. Introduction

The radio galaxy 3C326 N at $z=0.1$ is an excellent target to investigate how the mechanical energy output of radio-loud AGN affects the surrounding gas and star formation. In spite of a few $\times 10^9 M_\odot$ of dense H_2 , akin to LIRGs, the star-formation rate ($\text{SFR} \leq 0.07 M_\odot \text{ yr}^{-1}$ Ogle et al. 2007) is orders of magnitude lower than expected from the Schmidt-Kennicutt law (Nesvadba et al. 2010, N10 hereafter). With the low SFR, an X-ray faint AGN ($\log \mathcal{L}_X = 40.6 \text{ erg s}^{-1}$ Ogle et al. 2010), and a strong radio source (kinetic power $\log \mathcal{L}_{\text{kin}} \geq 44.6 \text{ erg s}^{-1}$, N10) 3C326 N provides a rare opportunity to clearly disentangle the effects of the jet from those of star formation and AGN radiation. It is one of the largest and oldest radio galaxies on the sky, with 2 Mpc size, dynamical and spectral ages of $6\text{--}20 \times 10^7 \text{ yrs}$ (Willis & Strom 1978) and a radio core at 3 mm (N10). 3C326 N is not an obvious member of a group or cluster, but has a similarly old ($\geq 10 \text{ Gyr}$), massive ($M_{\text{stellar}} = \text{few} \times 10^{11} M_\odot$) companion at a projected distance of $\sim 20 \text{ kpc}$.

3C326 N is amongst the 30% of nearby 3CR radio galaxies which have bright line emission from warm molecular hydrogen observed with Spitzer/IRS and line ratios consistent with shocks (Ogle et al. 2010). The gas kinetic energy and line luminosities exceed the energy injection rates from star formation and AGN radiation, making the deposition of mechanical energy by the radio source the only plausible culprit (N10). This energy injection does not only trigger an outflow, but also heats the ambient

gas through shocks, where the turbulent kinetic energy roughly equals the bulk kinetic energy of the outflowing gas. About half the molecular gas in 3C326 N is warm ($T > 100 \text{ K}$), a fraction $100\times$ greater than in star-forming galaxies, and suggesting that mechanical heating by the jet could be an important mode of suppressing the overall star formation in this galaxy.

We present deep imaging spectroscopy of the ro-vibrational H_2 emission lines at $R=3000$ and $0.7''$ spatial resolution, probing the spatially resolved gas properties and kinematics, which significantly enhances and complements the analysis of N10 which relied on unresolved Spitzer/IRS spectroscopy at $R=120$ (2500 km s^{-1}) with $5''$ spatial resolution. We find a kpc-sized rotating, turbulent disk with broad complex line profiles, suggesting that the mechanical energy deposition of the jet is not restricted to a small part of the disk. We argue that finding such a disk is difficult to reconcile with the recent star-formation history of 3C326 N unless the molecular gas formed in response to the turbulence created by the radio jet, indicating a close symbiosis between jet and molecular gas, but also suggesting that boosting the formation of molecular gas in turbulent jet cocoons is not sufficient to trigger star formation as often assumed in models of positive AGN feedback. Throughout the paper we adopt a $H_0 = 70 \text{ km s}^{-1}$, $\Omega_M = 0.3$, $\Omega_\Lambda = 0.7$ cosmology.

2. The molecular gas in 3C326 N

Data were obtained with the near-infrared imaging spectrograph SINFONI (Bonnet et al. 2004) at the Very Large Telescope of ESO with a total of 15,000 seconds of on-source observing time in the K-band at a seeing-limited resolution of

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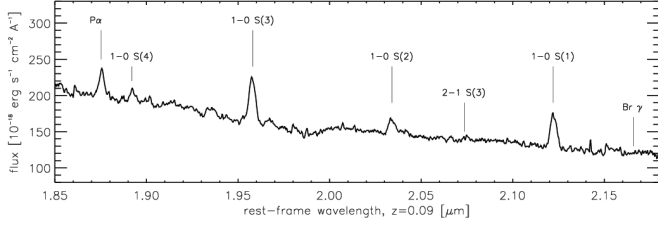


Fig. 1. Integrated K-band spectrum of 3C326 N. The properties of labeled lines are listed in Table 1.

$\text{FWHM}=0.75'' \times 0.65''$. Our data reduction and calibration methods are described by, e.g. Nesvadba et al. (2008, 2011a,b). The integrated spectrum of 3C326 N (Fig. 1) is dominated by the rovibrational H_2 lines H_2 1-0 S(1) to S(4) (S(1)-S(4) hereafter). $\text{Pa}\alpha$ is very weak in comparison, $\text{Pa}\alpha/\text{S}(3)=0.5$ (Table 1).

Several arguments suggest this gas is heated by shocks, similar to the pure-rotational H_2 and optical lines (N10). First, $\text{S}(3)/\text{Pa}\alpha$ and $\text{S}(1)/\text{Br}\gamma$ ratios are factors 10 greater than the typical $\text{S}(1)/\text{Br}\gamma} \sim 0.1 - 1.5$ in star-forming galaxies dominated by UV heating (Puxley et al. 1990). Since we do not detect $\text{Br}\gamma$, we estimate $\text{S}(1)/\text{Br}\gamma}=18-22$ from the $\text{Pa}\alpha$ flux and a line decrement $\text{Pa}\alpha/\text{Br}\gamma}=12-14$, which exceeds ratios in star-forming galaxies by a factor 10. Second, for star formation we expect ratios of the H_2 2-1/ H_2 1-0 S(3) lines factors of a few greater than the 3σ upper limit of 0.12 for S(1) and S(3) (Le Petit et al. 2006, see Herrera et al. 2011 for a similar analysis). A third diagnostics is the $[\text{OI}]\lambda 6300/\text{H}\alpha$ ratio as a function of $\text{S}(1)/\text{Br}\gamma$ (Mouri et al. 1989). For the $[\text{OI}]\lambda 6300/\text{H}\alpha$ ratio of N10 the $\text{S}(1)/\text{Br}\gamma$ ratio in 3C326 N exceeds that expected from star formation by a factor 10. Forth, the H_2 line luminosity of S(1)-S(3), $1 \times 10^{41} \text{ erg s}^{-1}$, exceeds the X-ray luminosity of the AGN and the mechanical luminosity of star formation (N10), leaving shock heating by the jet as the only plausible culprit.

$\text{Pa}\alpha$ and H_2 morphologies are disk-like and extend over $2.0'' \times 1.1''$ ($3.2 \text{ kpc} \times 1.8 \text{ kpc}$) at $\text{PA}_{\text{morph}}=0 \pm 20^\circ$. All H_2 lines and $\text{Pa}\alpha$ are broad, $\text{FWHM} \sim 400-700 \text{ km s}^{-1}$, down to the limit of our spatial resolution. They have irregular profiles (Fig. 2) that cannot be fit with single Gaussians. To map the kinematics, we therefore measure the width, W_p , and central wavelength, C_p , of each line at $p=0.2, 0.5, 0.7$, and $0.9 \times$ the line core extracted from small apertures of $0.4'' \times 0.4''$ (Heckman et al. 1981). W_{50} corresponds to the FWHM, and C_{90} approximates the line core. For S(3) we also measure the asymmetry parameter A_{20} , i.e., the shift of C_{20} compared to C_{90} (Fig. 2 of Heckman et al. 1981). S(1) and $\text{Pa}\alpha$ have telluric absorption features along their wings which makes measuring A_{20} difficult, but their C_{90} and W_{50} give results consistent with S(3).

Velocities increase monotonically from SSW to NNE along $\text{PA}=345^\circ \pm 15^\circ$ (Fig. 3) common to all percentiles and all lines. Lines are broadest NE of the nucleus ($W_{50, \text{H}_2} \leq 650 \text{ km s}^{-1}$, $W_{50, \text{Pa}} \leq 750 \text{ km s}^{-1}$), and more narrow in the south ($300-400 \text{ km s}^{-1}$ for all lines). A Fourier analysis of non-circularities in the velocity map with Kinemetry (Krajinović et al. 2006) suggests the velocity gradient is indistinguishable from rotation with $\text{PA}_{\text{kin}}=340^\circ$, inclination $i=35^\circ$ and a deprojected circular velocity of $v_c = 1/2 \Delta v_{\text{proj}} \csc i = 290 \text{ km s}^{-1}$. This would imply a dynamical mass $M_{\text{dyn}} = v_c^2 R/G = 3 \times 10^{10} M_\odot$ somewhat larger than the enclosed stellar mass of $M_{\text{stel}} = 2 \times 10^{10} M_\odot$. (The enclosed stellar mass is derived from the total stellar mass of $3 \times 10^{11} M_\odot$ (N10) for a mass profile consistent with a Hubble

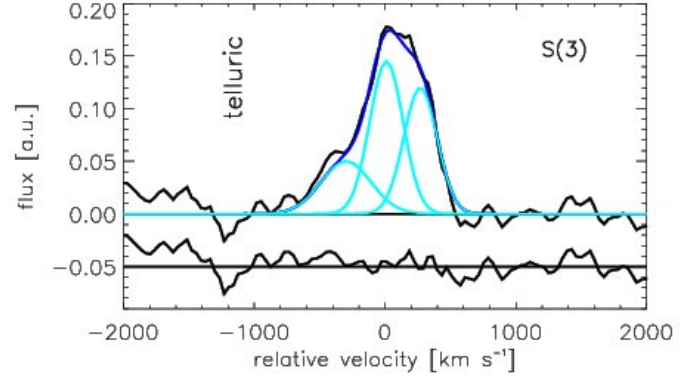


Fig. 2. Emission-line profile of H_2 1-0 S(3) extracted from a $0.6'' \times 0.6''$ aperture on the Northern disk (black line) fit with 3 Gaussian components (light blue line). The fit residual is shown below.

law.) The non-Gaussian line profiles suggest non-circular velocity components which are probably smoothed out by the low spatial resolution. This could affect the velocity gradient.

The FWHMs are larger than the velocity gradient, hence the broad lines cannot be artifacts of the velocity gradient and beam smearing, in which case most of the broadening should be near the nucleus, contrary to Fig. 3. The analysis of a toy data cube with constant $\text{FWHM}=400 \text{ km s}^{-1}$ and a gradient as observed suggests that beam smearing accounts for $\leq 15\%$ of the FWHM (similar to our 1σ observational errors). We also used a toy data cube to quantify the contribution of rotation to the CO(1-0) line measured by N10 which has $\text{FWHM}=350 \text{ km s}^{-1}$. We find a minimal intrinsic line width of 250 km s^{-1} assuming that the CO has the same velocity gradient as H_2 and $\text{Pa}\alpha$ and uniform surface brightness.

The small ratios of $v_c/\sigma = 1.7 - 0.9$ (with velocity dispersion $\sigma = \text{FWHM}/2.355 = 170 - 270 \text{ km s}^{-1}$, Fig. 3) imply an ellipsoidal configuration rather than a thin disk. Broad lines across all of the disk show that the mechanical energy of the AGN affects the gas globally. This differs strongly from the scenario of jet-cloud interactions which are confined to gas in small areas along the jet axis and is more akin to recent hydrodynamic models of jet cocoons where radio-emitting plasma permeates inhomogeneous gas disks along relatively low-density channels (Sutherland & Bicknell 2007; Wagner & Bicknell 2011).

The H_2 emission lines in 3C326 N have blue asymmetries which are strongest in the north ($A_{20} = -150 \text{ km s}^{-1}$, where turbulent velocities are also largest). H_2 emission has velocities of up to -1000 km s^{-1} from C_{90} , well above the local escape velocity, $v_{\text{esc}} \sim \sqrt{2} v_c \sim 400 \text{ km s}^{-1}$ at the largest H_2 radii. v_{esc} is also very similar to the lowest FWHMs in the South. Fig. 2 suggests that of-order 10% of the emission is from gas at $v > v_c$. Since the V(1-0) H_2 lines only probe the warmest, but not the bulk of the gas, this is difficult to turn into a mass outflow rate \dot{M} . If we assume that the pure-rotational lines, which trace most of the mass, have a similar profile (as recently reported for AGN with high-resolution Spitzer spectroscopy, Dasyra & Combes 2011, Guillard et al. 2011 ApJ submitted), we find $\sim 1 \times 10^8 M_\odot$ of entrained dense, warm H_2 in the wind. For an outflow timescale of 10^{7-8} yrs, this is consistent with or less than \dot{M} estimated from Na D, $30 - 40 M_\odot \text{ yr}^{-1}$ (N10), which traces neutral (atomic and/or molecular) gas down to lower column densities.

3. Jet-disk symbiosis

The gas in 3C326 N is in a kpc-sized, dense, thick disk. The broad line widths and complex profiles are not consistent with pure rotation, but require an additional source of turbulence which, following the energy arguments of N10, can only be the radio source. A fraction of the line emission is from gas at velocities above the local escape velocity. The rotational time of this disk, $\tau_{rot} = 2\pi R/v_c = 3 \times 10^7$ yrs, is shorter than the 6-20 $\times 10^7$ yrs age of the radio source, and comparable to the depletion time of the gas through the wind, $\tau_{wind} = M/\dot{M} \sim 5 \times 10^7$ yrs. It is also very similar to the dissipation time of turbulent energy $\tau_{diss} = R_{disk}/\sigma_{CO} = 1.5 \times 10^7$ yrs. The similar time scales raise two major questions.

First, why do we find a fairly organized (albeit thick) disk, if the depletion time through the wind is short? The mass outflow rate of N10 may be uncertain to about an order of magnitude, but even if it were a factor 10 lower, much of the gas should have been removed in the 2×10^8 yrs time span suggested by the oldest radio emission (§1). Finding a disk associated with an old radio source like 3C326 requires that energy losses are roughly balanced by the injection of energy, and the entrained by cooling gas. This would imply that a fraction of the outflowing gas does not escape from the galaxy, but dissipates its kinetic energy, stalls and rains back onto the disk in a cycle that may be akin to ‘cold feedback’ models of galaxy clusters (Ciotti et al. 1991; Pizzolato & Soker 2005; Brighenti & Mathews 2006).

Second, with $\tau_{diss} \sim \tau_{rot}$ the H_2 disk cannot be stable in absence of the jet. The absence of a gas-rich merger suggests the material accumulated gradually from mass return and accretion over $\gtrsim \text{few} \times 10^9$ yrs (N10), longer than the age of the radio jet. Collisions between clouds would rapidly dissipate the turbulent kinetic energy and the angular momentum of the disk in about one to a few rotational times (Verdoes Kleijn et al. 2006; Pizzolato & Soker 2010), consistent with our estimates. For the measured gas surface density $\Sigma_{gas} = 250 M_\odot \text{ pc}^{-2}$, the Schmidt-Kennicutt law implies $\text{SFR} = 4.5 M_\odot \text{ yr}^{-1}$. This is $\sim 60\times$ greater than the upper limit of $\text{SFR} = 0.07 M_\odot \text{ yr}^{-1}$ (Ogle et al. 2007) from Spitzer 70 μm imaging¹, and at odds with the recent star-formation history of 3C326 N. To quantify recent star formation we repeated the population synthesis analysis of the inner 3'' of 3C326 N from N10, this time adding small fractions of light from young ($3\text{--}30 \times 10^7$ yrs) stellar populations (YSPs) to the data. We find that at $\text{SFR} = 4.5 M_\odot \text{ yr}^{-1}$ stars must form for a few $\times 10^5$ yrs to produce detectable signatures. This is very short, hence, the absence of recent star formation suggests that the gas cannot have lost turbulent support for much longer than a free-fall time (10^6 yrs for $N = 10^3 \text{ cm}^{-3}$) in the last 300 Myr. Unless the jet has been nearly continuously active during this time (longer than the oldest estimate of the jet age and unlike many other giant radio galaxies with signs of repeated, shorter jet outbreaks, Schoenmakers et al. 2000), then why has this gas not been forming stars?

Resolving this paradox is possible if the molecules formed *after* the onset of AGN activity, boosted by the turbulence in the jet cocoon. Hydrodynamic simulations suggest that turbulent compression can turn a diffuse inhomogeneous, largely atomic medium within few 10^6 yrs into a warm, largely molecular medium (Glover & Mac Low 2007). Molecules form in rather dense regions ($n \geq 1000 \text{ cm}^{-3}$) and are transported to regions of lower density. The total amount of H_2 on macro-scales is set by the dynamic equilibrium of the rapid molecule formation and destruction on micro-scales. This gas would not be gravitation-

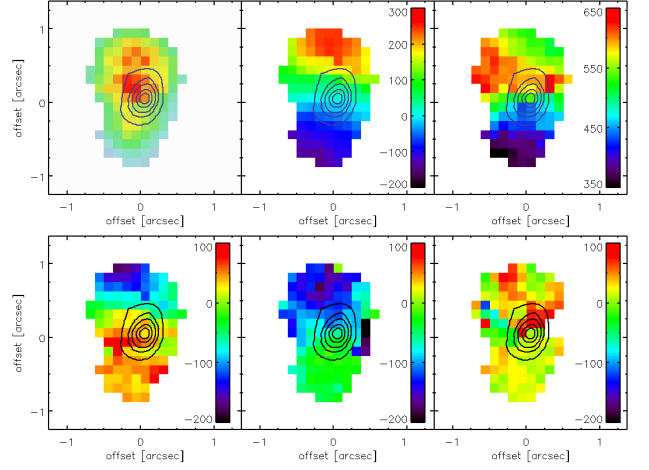


Fig. 3. (top, left to right) Maps of the line fluxes, relative velocities, and FWHMs of H_2 1-0 S(3) in km s^{-1} . (bottom, left to right) Line asymmetry of the 20-percentile for S(1), S(3), and $\text{Pa}\alpha$, respectively, in km s^{-1} . Contours mark the continuum peak.

ally bound and be part of a turbulent multiphase medium, where warm H_2 becomes a major gas coolant (Guillard et al. 2009).

In this scenario the molecular gas only survives the radio phase if it forms self-gravitating clouds while pressurized by the cocoon. Otherwise, with decreasing pressure H_2 formation rates drop, and the dynamic equilibrium between molecule formation and destruction is no longer maintained. The gas becomes again atomic. Extended reservoirs of diffuse atomic gas are not unusual in early-type galaxies (Oosterloo et al. 2007). Dust lanes in radio galaxies are often perpendicular to the jet (de Koff et al. 2000), and have masses which depend on jet power (de Ruiter et al. 2002), suggesting the ISM ‘‘knows about the jet’’.

4. Turbulence in 3C326 N and star formation

To better understand how AGN and star formation in 3C326 N may be related, we will now argue that turbulence governs the gas dynamics in 3C326 N down to sizes of molecular clouds. In this case finding low star formation rates as observed may not be so surprising after all. We quantify the ratio of turbulent to gravitational energy through the virial parameter, $\alpha_{vir} = 5\sigma_{cl}^2(\pi G R_{cl} \Sigma_{gas})^{-1}$ (Bertoldi & McKee 1992), where Σ_{cl} and R_{cl} are the mass surface density and radius of individual clouds, respectively, and G is the gravitational constant. Gravitationally bound clouds have $\alpha_{vir} \leq 1$. Hydrodynamic simulations suggest that column densities created by turbulent motion cannot greatly exceed the mean column density (e.g., Ostriker et al. 2001). We therefore use the measured $\Sigma_{gas} = 250 M_\odot \text{ pc}^{-2}$ (§2).

The H_2 line emission in 3C326 N is powered by the dissipation of turbulent energy, which implies $L_{H_2}/M_{H_2} = 3/2 f_{H_2} \sigma_{cl}^3/R_{cl}$ (McKee & Ostriker 2007), where L_{H_2} and M_{H_2} are the H_2 luminosity and mass. N10 find $L_{H_2}/M_{H_2} \sim 0.06 L_\odot/M_\odot$. The correction factor $f_{H_2} \lesssim 1$ is necessary because H_2 is the dominant, but not the only gas coolant (see also Herrera et al. 2011, A&A accepted). We adopt $f_{H_2} = 0.5$. Turbulent velocities do not scale arbitrarily with size, but approximately follow a power law $\sigma \propto R^{0.5}$, as observed for molecular clouds in the Milky Way (e.g., Larson 1981), and seen in hydrodynamic simulations including simulations of jet cocoons (Krause & Alexander 2007).

Fig. 4 illustrates that for all cloud sizes $\lesssim 1$ kpc, the observed L/M implies line widths at least $10\times$ greater than expected for

¹ for $R=1.6$ kpc. N10 found a smaller offset with a fiducial $R=2.5$ kpc

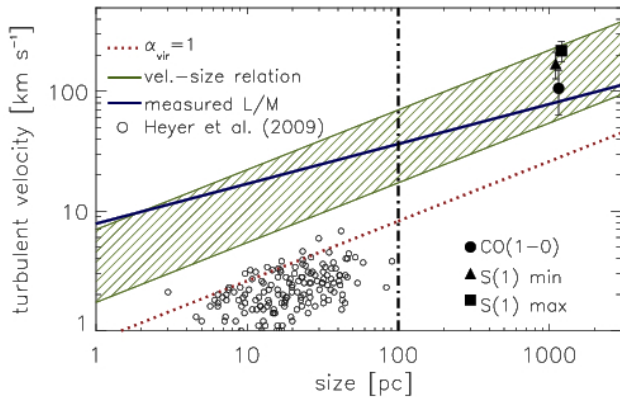


Fig. 4. Turbulent velocity as a function of size. Black symbols show our data for CO, H₂ 1-0 lines, and Pa α , respectively. The red dotted line is for $\alpha_{\text{vir}} = 1$ for a constant mass-surface density $\Sigma_{\text{gas}} = 250 \text{ M}_{\odot} \text{ pc}^{-2}$ and the green hatched region shows the range implied by Larson's scaling law and the range of velocities measured on kpc scales. Black circles show the molecular clouds of Heyer et al. (2009). These points fall near the $\alpha_{\text{vir}} = 1$ relation for gas surface densities ~ 5 smaller than that of 3C326 N. The horizontal dot-dashed line shows the maximal clump size of ~ 100 pc. See §4 for details.

$\alpha_{\text{vir}} = 1$ (red line in Fig. 4), suggesting that the turbulent kinetic energy strongly exceeds the gravitational energy on all relevant scales. This includes for example the ~ 100 pc scale at which disks akin to that of 3C326 N, but without additional turbulence from the jet, would fragment (Escala & Larson 2008, dot-dashed vertical line in Fig. 4), and also the cloud sizes of few 10s pc typically adopted in simulations of turbulent jet cocoons (Wagner & Bicknell 2011; Antonuccio-Delogu & Silk 2008).

Fig. 4 shows also that turbulent velocities in 3C326 N are likely much greater than those in typical molecular clouds in the Milky Way (Heyer et al. 2009, grey circles in Fig. 4). This may hold the key to the low star-formation rates. Local density enhancements owing to turbulent compression may enable star formation in clouds with $\alpha_{\text{vir}} \gtrsim 1$, but the star formation efficiency decreases rapidly with increasing α_{vir} (e.g., Krumholz & McKee 2005; Padoan & Nordlund 2011).

These arguments rely on the assumption that Larson's scaling relationship approximately holds for 3C326 N. Resolving scales of ~ 100 pc will only become possible once ALMA reaches its full capabilities, however, we already see in Fig. 4 that Larson's scaling law (green hatched area) normalized to the range of velocities at kpc scales falls within the range of turbulence required to account for the observed H₂ luminosity.

5. Concluding remarks

Our results suggest that AGN can create large reservoirs of dense, rapidly cooling molecular gas, as is suggested by models of AGN-triggered star formation (e.g., Mellema et al. 2002; Fragile et al. 2004; Silk & Norman 2009), but this does not *per se* seem sufficient to form stars. The turbulent energy injected by the jet seems capable to keep significant fractions of the gas warm ($T > 150$ K) and gravitationally unbound, making star formation very inefficient.

3C326 N is certainly an extreme example of a gas-rich radio galaxy with a very low star-formation efficiency, however, it is not unique. 30% of nearby 3CR radio galaxies have bright line

emission from warm H₂ (Ogle et al. 2010), and emission-line diagnostics suggest that much of this gas could be heated mechanically through shocks (N10). This includes galaxies like 3C293, which have on-going star formation, albeit not at the level expected from their H₂ reservoirs (Papadopoulos et al. 2010). It will be interesting to investigate in more depth how star formation and molecular gas kinematics are related in these cases.

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Line	λ_0	λ_{obs}	FWHM	flux
	[μm]	[μm]	[km s^{-1}]	[$10^{-15} \text{ erg s}^{-1} \text{ cm}^{-2}$]
Pa α	1.875	2.0443 ± 0.0004	516 ± 44	1.2 ± 0.14
H ₂ 1-0 S(1)	2.122	2.3132 ± 0.0005	534 ± 32	1.8 ± 0.13
H ₂ 1-0 S(2)	2.034	2.2170 ± 0.0003	563 ± 45	0.7 ± 0.05
H ₂ 1-0 S(3)	1.958	2.1339 ± 0.0004	594 ± 41	2.2 ± 0.2
H ₂ 1-0 S(4)	1.892	2.0624 ± 0.0004	475 ± 45	0.8 ± 0.05
Br γ	2.166			$< 0.26 (3\sigma)$
H ₂ 2-1 S(3)	2.074			$< 0.24 (3\sigma)$

Table 1. Line properties derived from the “stacked” spectrum, i.e., the integrated spectrum with the large-scale velocity gradient removed. We also list the 3σ upper limit for Br γ , and H₂ 2-1 S(3), which are not detected.